

# Ultra-Heavy Cosmic Ray Analysis with CALET on the International Space Station: Established and Developing Procedures

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The CALorimetric Electron Telescope (CALET) has collected over 60 months of uninterrupted data on the flux and spectrum of the Ultra-Heavy (UH) cosmic rays from  $Z=30$  to 40. Using the latest data provided from CALET's UH trigger, we present a newly developed UH analysis complementary to the ongoing analysis presented at this conference by Zober et al. This work introduces a new Ultra-Heavy Analysis (UHA) dataset produced from CALET production data allowing for more streamlined analysis. We detail temporal and spatial correction algorithms using both the  $^{26}\text{Fe}$  and  $^{14}\text{Si}$  peaks to improve charge resolution in the  $Z \geq 30$  region. Additionally, this work presents a new method for removing the contributions from non-relativistic/lower- $Z$  nuclei using trajectory tracing to compute effective cutoff rigidities in place of the previously used vertical Stoermer approximation. We show that replacing the approximated cutoffs with numerically calculated effective cutoff rigidities, calculated using the IGRF13 and T05 (Tsyganenko 05) geomagnetic field models, leads to fewer events being removed from the dataset while maintaining improved charge resolution for  $Z > 26$ . Furthermore, we introduce Tarle function peak fitting to perform charge corrections needed as a result of any quenching effects. We show the most recent CALET UH results incorporating these improvements in the analysis

*37<sup>th</sup> International Cosmic Ray Conference (ICRC 2021)*

*July 12th – 23rd, 2021*

*Online – Berlin, Germany*

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\*Presenter

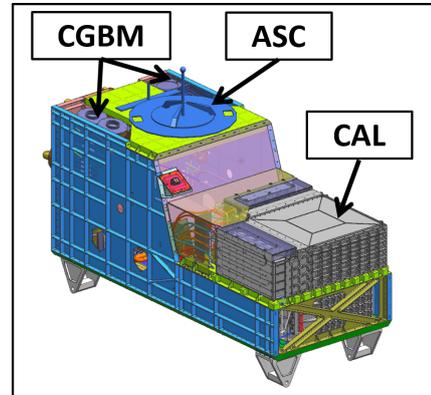
## 1. Introduction

The CALorimetric Electron Telescope (CALET) has collected over 60 months of uninterrupted data on the flux and spectrum of the Ultra-Heavy (UH) cosmic rays from  $Z=30$  to 40. Using the latest data provided from CALET's UH trigger, this work presents a newly developed UH analysis complementary to the ongoing analysis by fellow CALET collaborators. We detail temporal and spatial correction algorithms using both the  $^{26}\text{Fe}$  and  $^{14}\text{Si}$  peaks to improve charge resolution in the  $Z \geq 30$  region. Additionally, this work presents a new method for removing contamination from non-relativistic/lower- $Z$  nuclei using trajectory tracing to compute effective cutoff rigidities in place of the previously used vertical Stoermer approximation. We show that replacing the approximated cutoffs with numerically calculated effective cutoff rigidities, calculated using the IGRF13 and T05 (Tsyganenko 05) geomagnetic field models, leads to fewer events being removed from the dataset while maintaining improved charge resolution for  $Z > 26$ . Furthermore, we introduce Tarle function peak fitting to perform charge corrections needed as a result of any quenching effects. We show the most recent CALET UH results incorporating these improvements in the analysis.

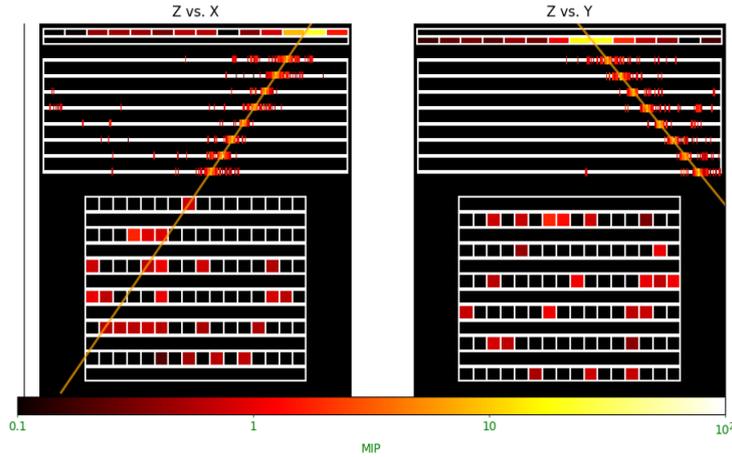
## 2. The CALET Instrument

CALET was launched in August 2015 to the International Space Station (ISS) for the primary purpose of directly measuring the cosmic-ray total electron spectrum up to energies of tens of TeV. The CALET calorimeter (CAL) [2], shown in Figure 1, has a depth of 30 radiation lengths ( $X_0$ ) at normal incidence and includes three separate sub-detectors. The Charge Detector (CHD) consists of two layers of 14 plastic scintillating paddles and can measure the charge of nuclei up to  $Z = 40$ . The Imaging Calorimeter (IMC) contains 16 layers of 448 scintillating fibers each which provide accurate tracking for most events. Most of the shower energy is deposited in the Total Absorption Calorimeter (TASC), with 12 layers of 16 PWO logs that can contain electromagnetic showers up to TeV energies.

For ultra-heavy cosmic ray observations CALET uses a dedicated trigger mode that is active throughout the ISS orbit [1]. Standard trigger modes require events pass through the TASC, limiting the field of view, but the UH trigger removes this restriction and only requires that events pass through the CHD and top half of the IMC [1]. This raises the geometric factor from  $\sim 1200\text{cm}^2\text{sr}$  to  $\sim 4000\text{cm}^2\text{sr}$ , and prevents UH event candidates similar to the one shown in Figure 2 from being removed from the dataset.



**Figure 1:** CALET structure indicating the CALET Gamma-ray Burst Monitor (CGBM), Advanced Stellar Compass (ASC), and calorimeter (CAL).



**Figure 2:** Ultra-heavy event candidate shown in the CALET event viewer. The calculated UH track is shown by the orange line.

### 3. Data Analysis

#### 3.1 Ultra-Heavy Analysis Dataset

To aid the UH data analysis a reduced high level analysis UH dataset was produced using 5 years of the CALET Level 2 Pass 4.1 dataset from October 2015 to September 2020. Included for each of the  $1.8 \times 10^8$  events is the associated summary information, energy-deposit sums for each layer of the instrument, orbital parameters of the ISS at the time of observation, UH track reconstruction parameters, and individual energy deposits for detector elements on or near the calculated track. The estimated charge of each event is calculated using the event's  $dE/dx$  and the Bethe equation, which reduces to Equation 1 assuming the particle is minimum ionizing.

$$\frac{dE}{dx} \propto z^2 \quad (1)$$

A selection is then made to the initial set of events requiring their reconstructed track's incidence angle be less than 60 degrees [3]. These high incidence angle events can have their charges over estimated resulting in a reduction of charge resolution in the UH region. A charge consistency cut [3] is applied to the remaining events removing events where the estimated charges in the x and y layers of the CHD (CHDX and CHDY) differ by greater than 1.5%. The final dataset used for this analysis includes  $5 \times 10^7$  events.

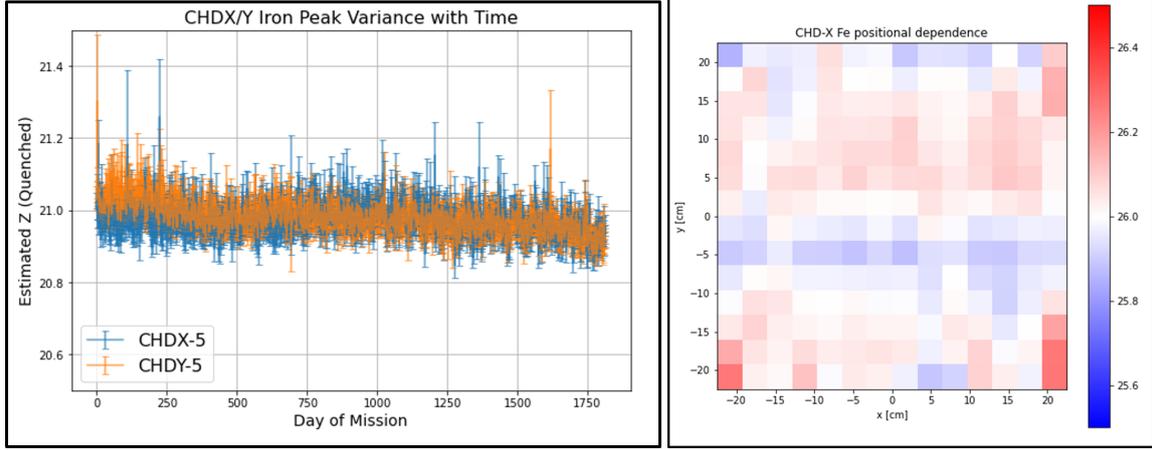
#### 3.2 Charge Corrections

##### 3.2.1 Time/Spatial Dependence

When calculating charge using the method described above, long-term changes in temperature and PMT gain cause time dependent fluctuations that negatively impact charge resolution. To address these variations, an example of which is shown in the left side of Figure 3, events are placed in 1-day bins and fits performed to find the location of the  $^{26}\text{Fe}$  and  $^{14}\text{Si}$  peaks for each paddle of both CHDX and CHDY. These peaks are then used to calculate linear correction parameters for

each day of data. When the corrections are applied, each bin's charge spectrum is linearly shifted so that the new  $^{26}\text{Fe}$  and  $^{14}\text{Si}$  peak locations are equal to their mean values for the full 5-year dataset.

Further variation in charge, shown in the right of Figure 3, arises from light propagation effects and non-uniformity in the CHD that both contribute to a spatial dependence in the light yield of the detector. To account for this dependence, the time-corrected data is divided into  $3.2\text{ cm} \times 3.2\text{ cm}$  bins in both the CHDX and CHDY. For each bin, the fitted  $^{26}\text{Fe}$  and  $^{14}\text{Si}$  peak locations are again used to calculate the linear correction needed to shift each peak to its mean value for the entire CHDX and CHDY respectively.



**Figure 3:** *left:* Location of the  $^{26}\text{Fe}$  peak for paddle 5 of both CHDX and CHDY as a function of time. *right:* Location of the  $^{26}\text{Fe}$  peak for each position bin of the CHDX.

### 3.2.2 Quenching Correction

In organic scintillators, it is found that large deposits of energy that are spatially concentrated will exhibit non-linear saturation of the scintillator response, or quenching. This results in underestimated charge assignments that worsens as  $Z$  increases. Using a method derived by Tarle et al., it is possible to recover the true  $dE/dx$  from the actual luminosity,  $dL/dx$ , via

$$\frac{dL}{dx} = K \frac{(1 - F_s)dE/dx}{1 + B(1 - F_s)dE/dx} + F_s \frac{dE}{dx}, \quad (2)$$

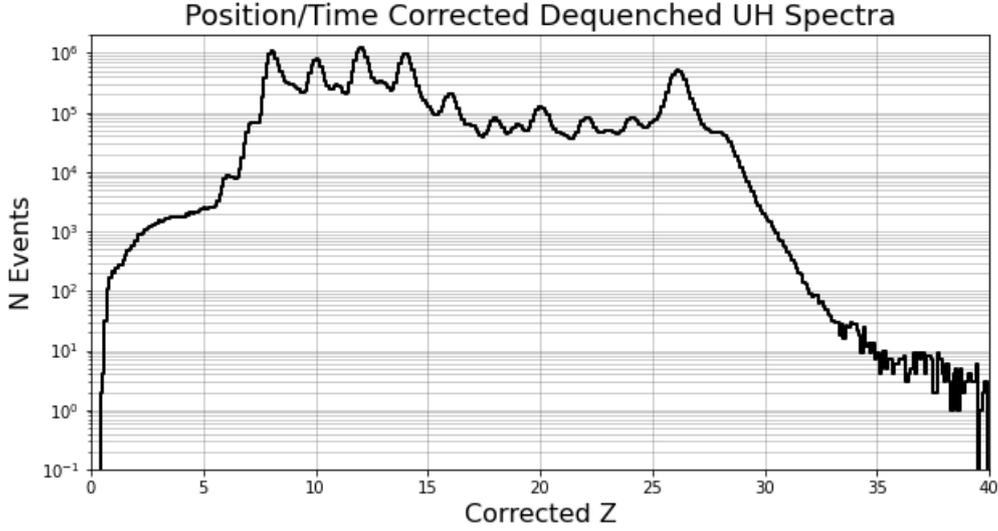
where  $F_s$  describes the fraction of electrons produced in the ionization escape between the saturated core and unsaturated halo regions of the signal, and  $B$  and  $K$  are constants. Combining Equations 1 and 2, we get

$$\frac{dL}{dx} = T(Z^2) = \frac{B_1 Z^2}{1 + B_2 Z^2} + B_3 Z^2. \quad (3)$$

Inverting this result then lets us solve for the un-quenched  $dE/dx$  directly from our measured luminosity,  $dL/dx$ , using

$$Z^2 = \frac{1}{2B_2 B_3} \left[ \left( \frac{dL}{dx} B_2 - B_3 - B_1 \right) + \sqrt{\left( \frac{dL}{dx} B_2 - B_3 - B_1 \right)^2 + 4 \frac{dL}{dx} B_2 B_3} \right], \quad (4)$$

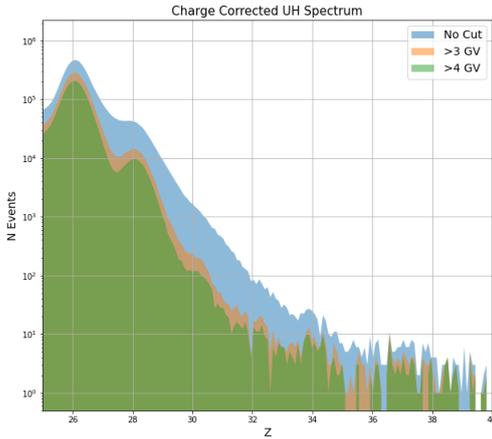
where  $B_1$ ,  $B_2$ , and  $B_3$  are fit constants. To calculate the necessary Tarle function parameters, we performed a multi-peak fit for all nuclei from  $_6$  to  $_{28}\text{Ni}$  and fit these directly with the Tarle function. The result of applying this quenching correction is shown in Figure 4.



**Figure 4:** Quenching corrected CHDX charge spectra for time and position corrected UH events.

### 3.3 Charge Resolution Improvements

#### 3.3.1 Stoermer Cutoff Rigidity



**Figure 5:** Charge corrected UH spectrum with no cut, a cut at  $R > 3$  GV, and a cut at  $R > 4$  GV. Where  $R$  is the vertical cutoff rigidity.

To achieve adequate charge resolution in the UH ( $Z \geq 30$ ) region, a method was required for removing the primary source of contamination, non-relativistic and lower  $Z$  nuclei that are below minimum ionizing and consequently have significantly increased  $\frac{dE}{dx}$ . Given CALET's orbit is deep within the Earth's geomagnetic field and experiences a wide range of cutoff rigidities, implementing a cut using geomagnetic cutoff rigidity was the chosen method. Non-relativistic nuclei below minimum ionization have access to the instrument through a smaller, lower range of rigidities than higher energy nuclei, so removing times where the cutoff rigidity is in this range limits their inclusion in the dataset.

Initially, Stoermer's vertical cutoff rigidity approximation [4], Equation 5, was used for determining cutoffs where  $\Psi$  is magnetic latitude.

$$R_c = 14.9 * \cos^4(\Psi), \quad (5)$$

Figure 5 shows the results of implementing a cut using this approximated cutoff, showing cut thresholds of 3 GV and 4 GV. Although there is visible improvement in the charge resolution, the loss of statistics in the UH region limits quickly becomes prohibitive to the analysis.

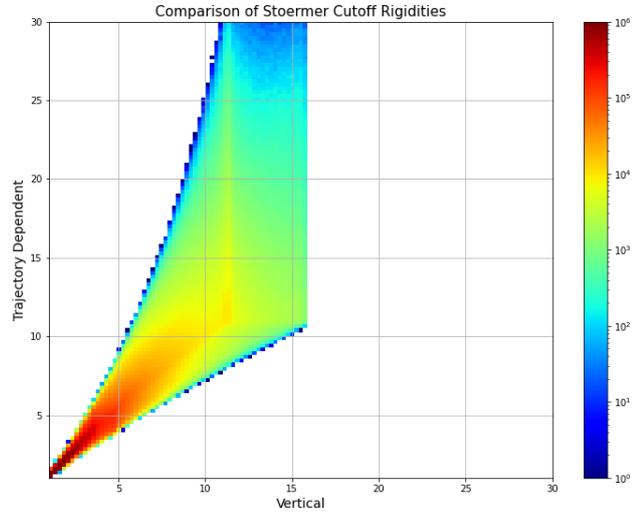
To refine the calculation of the cutoff rigidities, the vertical Stoermer approximation was replaced with its trajectory dependent equivalent [4]. Instead of assuming normal incidence, Equation 6 now takes into account the incident particle's zenith,  $\phi$ , and azimuth,  $\theta$ , angles, as well as its distance from the magnetic dipole,  $r$ , when deriving the cutoff. A comparison between the computed cutoff rigidities is shown in Figure 6, and shows significant divergence between the vertical and trajectory dependent values.

$$R_c = \frac{60}{r^2} \times \left( \frac{1 - \sqrt{1 - \cos(\Psi)\cos^3(\theta)}}{\cos(\Psi)\cos(\theta)} \right)^2 \quad (6)$$

This divergence though appears primarily at times of high cutoff rigidity, and does not affect events observed over much of the orbit.

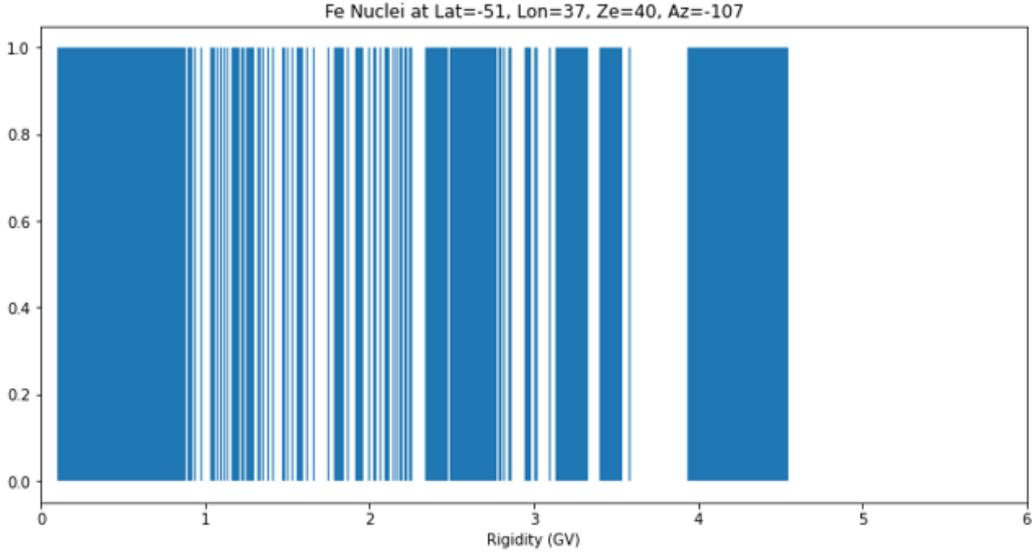
### 3.3.2 Numerical Method for Calculating Cutoff

To improve on the Stoermer approximation, a trajectory tracing method developed by Smart and Shea [5] was implemented to numerically solve for effective cutoff rigidities on an event by event basis. For each trajectory, the method uses an adaptive step size Runge-Kutta method to integrate the equations of motion for a charged particle in a magnetic field, stepping through the trajectory until it is deemed successful or forbidden. This is repeated for a range of rigidity values for each event, resulting in a spectrum of allowed and forbidden rigidities. An example of this, shown in Figure 7, shows all allowed rigidities in white and forbidden rigidities in blue. Notable points are the upper cutoff rigidity where the first forbidden trajectory occurs, and the lower rigidity where the last allowed trajectory occurs. The effective cutoff rigidity is calculated using



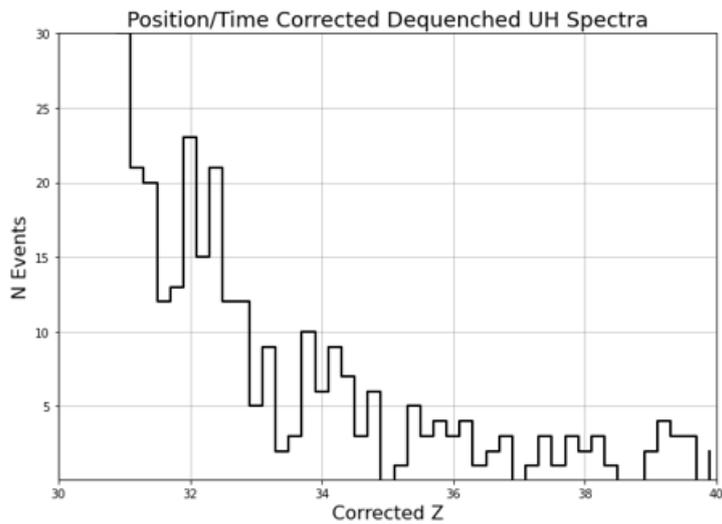
**Figure 6:** Comparison between trajectory dependent and vertical Stoermer cutoff rigidities.

$$R_E = R_L + (R_U - R_L) \times N_{Forbidden}/N_{Total}, \quad (7)$$



**Figure 7:** Spectrum of allowed and forbidden trajectories for an example  ${}_{26}\text{Fe}$  event trajectory, with allowed trajectories being white space, and forbidden trajectories being blue space.

where  $R_L$  is the lower cutoff,  $R_U$  is the upper cutoff,  $N_{Forbidden}$  is the number of forbidden trajectories after  $R_U$ , and  $N_{Total}$  is the total number of trajectories between  $R_U$  and  $R_L$ . This effective rigidity is the value used for the calculated cutoff rigidity, as it takes into account the effect of the penumbra (allowed trajectories interspersed amongst forbidden trajectories) and is a more appropriate value when compared to the upper or lower cutoff. The effective cutoffs were calculated for all events in the UH dataset using the IGRF13 and Tsyganenko 05 internal/external geomagnetic field models. The resulting UH spectrum, Figure 8, incorporating all the corrections discussed above, and using the effective cutoff rigidity cut, shows a clear improvement in charge resolution in the UH region.



**Figure 8:** UH charge spectrum with effective cutoff rigidities  $> 4$  GV.

#### 4. Conclusions

The analysis improvements described in this work have provided results consistent with previous CALET UH results, producing a UH spectrum with resolved peaks for evenly charged nuclei up to  $Z=38$ . Using numerically computed effective cutoff rigidities to limit contamination in the UH dataset has improved charge resolution in the UH region. Further refinements to the analysis, including analysis presented by Zober et al. [6], along with increased statistics from the extended CALET mission will continue to improve CALET UH results going forward.

#### 5. Acknowledgements

The material contained in this document is based upon work supported by a National Aeronautics and Space Administration (NASA) grant or cooperative agreement. Any opinions, findings, conclusions, or recommendations expressed in this material are those of the author and do not necessarily reflect the views of NASA. This work is supported in Japan by JAXA, in Italy by ASI and in the U.S. by NASA.

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